

## Gauge independence of induced gravitational waves

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We study gauge (in)dependence of the gravitational waves (GWs) induced from curvature perturbations. For the GWs produced in a radiation-dominated era, we find that the observable (late-time) GWs in the transverse-traceless (synchronous) gauge and in the Newtonian gauge are the same in contrast to a claim in the literature. We also mention the interpretation of the gauge dependence of the tensor perturbations which appears in the context of the induced GWs.

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### I. INTRODUCTION

One of the main predictions of cosmological inflation is the generation of the cosmic perturbations. So far, there is only an upper bound on the amplitude of tensor perturbation, but that of scalar perturbations, namely the curvature perturbations, has been well established by the precise observations of the cosmic microwave background and the large-scale structure [1,2]. Although these observations provide us information on the large scale, we cannot access much smaller scales due to the Silk damping of the radiation perturbations or the nonlinear growth of the matter perturbations. Such information on the small-scale perturbations is desirable to understand the global picture of the inflaton potential and particle physics behind it.

Gravitational waves (GWs) are a useful probe of small-scale curvature perturbations [3–8] since if the latter is enhanced, a sizable amount of GWs is produced at the second order of the cosmological perturbation. We call it the induced GWs<sup>1</sup> (see early works [9–12] and recent developments [13–22]). Such an enhancement of the curvature perturbations are often considered in the context of the primordial black hole (PBH) [23–25] (see also Refs. [7,26–31]). PBHs have recently been studied by many authors mainly because they can be a dark matter (DM) candidate and may explain the binary black hole merger events observed by LIGO/Virgo [32]. If the PBHs explain the DM or LIGO/Virgo events, the induced GWs can be observed by the future projects, such as LISA [33] and SKA [34], and even if there are few PBHs in the Universe, the future observations of the induced GWs could

determine or constrain the amplitude of the small-scale perturbations.

In contrast to the linear perturbation theory, the second-order (induced) tensor perturbations are known to have gauge dependence. A large change of the power spectrum of the tensor perturbations induced in a matter-dominated (MD) era was reported in Ref. [35]. More aspects of the gauge dependence were studied in Refs. [36–40].

In the context of the tensor perturbations induced by scalar perturbations, the tensor perturbation can be divided into two parts. One is the freely propagating tensor perturbation, following the equation of motion without any source. This kind of tensor perturbations is widely known as *gravitational waves* and the time dependence can be written as  $h_{ij} \propto \sin(k\eta)/a$  or  $\cos(k\eta)/a$  ( $a$ : scale factor), as is well known. Although they couple with scalar perturbations at their production, they finally decouple from the scalar perturbations and freely propagate. Once they decouple from the scalar perturbation, they are independent of the scalar perturbations and therefore do not depend on the gauge. The other is the tensor perturbation coupling with the scalar perturbations, in which the freely propagating tensor perturbations are included only until they decouple from the scalar perturbations. Since this kind of tensor perturbation is controlled by the scalar perturbations, the time dependence of this tensor perturbation inherits those of the scalar perturbations. Note that the gauge dependence appears in this kind of tensor perturbation. This tensor perturbation is also often called a gravitational wave in many references, but in this paper, to discriminate these two kinds of tensor perturbations, we call only the freely propagating tensor perturbation a *gravitational wave*.

<sup>1</sup>It is also called the second-order GWs, secondary GWs, or scalar-induced GWs, etc.

Recently, it was claimed<sup>2</sup> in Ref. [41] that the power spectrum of the induced GWs (freely propagating tensor perturbations) calculated in the TT gauge (also known as synchronous gauge) is reduced by at least one order of magnitude compared to that calculated in the Newtonian gauge (also called the Poisson gauge, the longitudinal gauge, or the zero-shear gauge), commonly used in the literature for the GWs produced during a radiation-dominated (RD) era.

However, we do not expect that the induced GWs depend on the gauge as we mentioned above. The gauge independence of the induced GWs can also be expected from the coincidence of the GWs calculated in the Newtonian gauge and the uniform curvature gauge (also called the flat gauge), shown in Ref. [40] (see the case of  $w > 0$  in the reference). In the present paper, we revisit the calculation in Ref. [41] and find that the difference between the tensor perturbations induced during a RD era in the TT gauge and in the Newtonian gauge decreases on subhorizon scales, and therefore the late-time (i.e., observable) GWs are the same in both gauges, in contrast to Ref. [41]. We also mention the interpretation of the gauge dependence reported in Ref. [40] and an important role played by the diffusion damping effect at the end of this paper.

## II. INDUCED GRAVITATIONAL WAVES IN TT GAUGE

In this section, we calculate the induced GWs in the TT gauge and compare them with those in the Newtonian gauge known in the literature. We take the convention in Refs. [14,42]. The comparison of the conventions in the literature is given in Appendix A. To avoid the repetition of steps of derivation, we write only the essential equations for a review. Details of derivation can be found in Refs. [14,42] (see also Ref. [36] for early works on the formulation in the TT gauge).

The GW energy density is given by [43]

$$\rho_{\text{GW}} = \frac{1}{16a^2} \overline{\langle h'_{ij} h'_{ij} \rangle}, \quad (1)$$

<sup>2</sup>One of the main claims in Ref. [41] is that the description of the detection of GWs with a space-based interferometer is most straightforward in the transverse-traceless (TT) gauge, and they advocate the use of the TT gauge (if the strength of the induced GWs is gauge dependent at all). The argument is as follows. The size of the space-based interferometer is typically not sufficiently large compared to the wavelength of the GWs. In such a situation, one cannot take the proper detector frame (Fermi normal coordinates) for the whole experimental region. The completely relativistic treatment is required, and they give at least two reasons that motivate one to use the TT gauge. For one thing, one has to otherwise add the integral of the gravitational potential along the photon path to make the time shift (measured as the phase shift of the laser) gauge invariant in other gauges. Note that the result of the TT gauge reduces to the simplest case where the description in the proper detector frame is possible. For the other, the sensitivity curve of LISA is calculated in the TT gauge.

where  $h_{ij}$  is the freely propagating tensor perturbation, and the overline denotes the oscillation average. Thanks to the oscillation average, we may calculate the spatial derivatives  $h_{ij,k} h_{ij,k}$  instead of the time derivatives, which is justified on subhorizon scales.

The current value of the power spectrum of the induced GWs per a logarithmic wave number bin  $\Omega_{\text{GW}}(\eta, k) \equiv \rho_{\text{GW}}(\eta, k)/\rho_{\text{total}}$  with  $\rho_{\text{GW}}(\eta) = \int d \ln k \rho_{\text{GW}}(\eta, k)$  is obtained from

$$\Omega_{\text{GW}}(\eta_0, k) = 0.387 \Omega_r \left( \frac{g_*(T_c)}{106.75} \right) \left( \frac{106.75}{g_{*,s}(T_c)} \right)^{\frac{4}{3}} \Omega_{\text{GW}}(\eta_c, k), \quad (2)$$

where  $\eta = \eta_0$  is the current conformal time; a subscript c denotes a time at which the tensor perturbation is dominated by the freely propagating tensor perturbation<sup>3</sup>;  $g_*$  and  $g_{*,s}$  are the numbers of relativistic degrees of freedom for the energy and entropy density, respectively; and  $\Omega_r$  is the current energy density fraction of radiation. This formula assumes the production of the GWs in a RD era. The constant value of  $\Omega_{\text{GW}}$  can be calculated as

$$\begin{aligned} \Omega_{\text{GW}}(\eta_c, k) &= \frac{1}{6} \left( \frac{k}{\mathcal{H}} \right)^2 \int_0^\infty dv \\ &\times \int_{|1-v|}^{1+v} du \left( \frac{4v^2 - (1+v^2 - u^2)^2}{4uv} \right)^2 \\ &\times \overline{I^2(u, v, x_c)} \mathcal{P}_\zeta(uk) \mathcal{P}_\zeta(vk), \end{aligned} \quad (3)$$

where  $x \equiv k\eta$  is a dimensionless time,  $\mathcal{H} = (da/d\eta)/a$  is the conformal Hubble parameter,  $\mathcal{P}_\zeta(k)$  is the power spectrum of the primordial curvature perturbations, and the late time limit  $x \rightarrow \infty$  is practically used for evaluation at  $x_c = k\eta_c$ . The function  $I(u, v, x)$  is defined as

$$I(u, v, x) = \int_0^x d\bar{x} \frac{a(\bar{\eta})}{a(\eta)} k G_k(\eta, \bar{\eta}) f(u, v, \bar{x}), \quad (4)$$

where  $\bar{x} = k\bar{\eta}$ , and  $G_k(\eta, \bar{\eta})$  is the Green function for the tensor perturbations, which is given as  $kG_k(\eta, \bar{\eta}) = \sin(x - \bar{x})$  in the production in a RD era. In this expression,  $u$  and  $v$  are symmetric under their exchange.<sup>4</sup>

<sup>3</sup>If we define the parameter  $\Omega_{\text{GW}}$  as the energy density parameter of the total tensor perturbations, as in many references, the subscript c denotes a time at which  $\Omega_{\text{GW}}$  reaches the asymptotic constant value.

<sup>4</sup>The power spectrum of the induced GWs is nothing but a 2-point correlation function, which can be written as a 4-point function of scalar source modes. In the absence of non-Gaussianity, the 4-point function can be expressed in terms of two different combinations of 2-point functions (two ways of Wick contraction). From this and from the ( $u \leftrightarrow v$ ) symmetry of the rest of the integrand as well as the integration region, the remaining asymmetry under the exchange of  $u$  and  $v$ , if any, must vanish. Once this is understood, one can freely symmetrize  $f(u, v, x)$  or equivalently  $I(u, v, x)$  as a convention.

Now, the final missing piece in the above equation is the explicit definition of the source function  $f(u, v, \bar{x})$ , which is different between the TT gauge and the Newtonian (or some other) gauge. The difference arises only from this. It can be derived from the expression of the scalar-scalar source term in the equation of motion for the second-order tensor perturbation, Eq. (46) of Ref. [39]. In terms of the transfer functions  $T_\psi$  and  $T_\sigma$  of the scalar perturbation variable  $\psi$  and the shear potential  $\sigma$  in the TT gauge, it is given by

$$f(u, v, x) \equiv -\left(\frac{2}{3}\right)^2 \left( -T_\psi(ux)T_\psi(vx) + \frac{3\sqrt{3}v}{2u}T'_\psi(vx)T_\sigma(ux) + \frac{3\sqrt{3}u}{2v}T'_\psi(ux)T_\sigma(vx) + \frac{3}{2uv}T_\sigma(ux)T_\sigma(vx) - uvx^2T'_\psi(ux)T'_\psi(vx) \right), \quad (5)$$

$$T_\psi(x) = \frac{9 - 9\cos(x/\sqrt{3})}{x^2},$$

$$T_\sigma(x) = \frac{-3\sqrt{3} + 9\sin(x/\sqrt{3})}{x^2}, \quad (6)$$

where the prime denotes the differentiation with respect to the argument, e.g.,  $T'(ux) \equiv \frac{dT(ux)}{d(ux)}$ . More explanations on the transfer functions and the choice of an integration constant are given in Appendix A. A difference of an overall factor from Ref. [41] is just because of different conventions. We find the coefficient of the first term inside the parenthesis in Eq. (5)  $-1$ , but  $2$  is used in Ref. [41].

$$\lim_{x \rightarrow \infty} \overline{\mathcal{I}^2(u, v, x)} = \frac{1}{2} \left( \frac{3(u^2 + v^2 - 3)}{4u^3v^3} \right)^2 \left( \left( -4uv + (u^2 + v^2 - 3) \log \left| \frac{3 - (u+v)^2}{3 - (u-v)^2} \right| \right)^2 + \pi^2 (u^2 + v^2 - 3)^2 \Theta(u + v - \sqrt{3}) \right), \quad (8)$$

in both the TT gauge and the Newtonian gauge, where  $\Theta(\cdot)$  is the Heaviside step function. Thus, once the power spectrum of the primordial curvature perturbations  $\mathcal{P}_\zeta(k)$  is specified, one can use Eqs. (2), (3), and (8) to compute the power spectrum of the induced GWs, and the result is irrespective of the choice of the TT gauge or the Newtonian gauge. Explicitly,

$$\Omega_{\text{GW}}^{(\text{TT gauge})}(\eta_0, k) = \Omega_{\text{GW}}^{(\text{Newtonian gauge})}(\eta_0, k), \quad (9)$$

and an analogous formula with  $\eta_0$  replaced by  $\eta_c$  hold.

In the remainder of this section, let us study the behavior of  $\mathcal{I}(u, v, x)$  in each gauge and their differences for finite  $x$ . Examples of the time evolution of  $\mathcal{I}(u, v, x)$  for given sets of  $u$  and  $v$  in both gauges are shown in Fig. 1. In the figure,  $x = k\eta$  can be interpreted as the time parameter in the unit in which the horizon entry of the

Also, our definition of the transfer functions is 3 times larger than that in Ref. [41] for the transfer function of the gauge invariant Bardeen potential  $\Psi$  to be unity for  $\eta \rightarrow 0$ . Therefore, we expect an amplitude and shape of the power spectrum of the induced GWs different from those of Ref. [41].

Similarly to the case of the Newtonian gauge studied in Refs. [13,14], we can analytically integrate Eq. (4) using the method of Ref. [10]. The result is long, so it is written in Appendix B. First of all, it is different from the case of the Newtonian gauge for finite  $x = k\eta$ . The difference between the two gauges is given by

$$\begin{aligned} & \mathcal{I}^{(\text{Newtonian gauge})}(u, v, x) - \mathcal{I}^{(\text{TT gauge})}(u, v, x) \\ &= \frac{18}{u^2v^2x} \left( 1 - \frac{\sqrt{3}}{ux} \sin \frac{ux}{\sqrt{3}} \right) \left( 1 - \frac{\sqrt{3}}{vx} \sin \frac{vx}{\sqrt{3}} \right) \\ &\simeq \min \left[ \frac{x^3}{18}, x, \frac{18}{u^2v^2x} \right], \end{aligned} \quad (7)$$

where  $\mathcal{I}(u, v, x) \equiv xI(u, v, x)$ , and the superscripts denote in which gauge it is evaluated. We introduce this combination  $\mathcal{I}$  as its oscillation average asymptotes to a constant while  $I$  decreases by redshift. In the second equality, we have used the fact that  $\max[u^2, v^2] \simeq 1$  when  $\min[u, v] \ll 1$ . From the above expression, it is clear that the gauge difference decreases in time on subhorizon scales, and the asymptotic constant values, which are relevant for direct observations, coincide with each other. That is,  $\lim_{x \rightarrow \infty} \mathcal{I}^{(\text{TT gauge})}(u, v, x) = \lim_{x \rightarrow \infty} \mathcal{I}^{(\text{Newtonian gauge})}(u, v, x)$ . Therefore, we can use the same formula [14] (see also Ref. [13])

tensor perturbation ( $k = \mathcal{H}$ ) occurs at  $x = 1$ . The variables  $u$  and  $v$  are to be interpreted as the ratio of the length scales of the tensor perturbation and the typical or relevant scalar source. For example, if we consider the monochromatic power spectrum of the curvature perturbation at  $k_*$ , taken in Ref. [41], the relevant values of  $u$  and  $v$  are  $k_*/k$ .

Let us begin with the small  $x$  (superhorizon) limit. The leading order terms of  $I(u, v, x)$  in this limit are  $(2/9)x^2$  in the Newtonian gauge<sup>5</sup> and  $(1/6)x^2$  in the TT gauge. In terms of  $\mathcal{I}$ , the difference is thus  $(1/18)x^3$ .

After the horizon entry of the scale of the would-be GWs, the difference in the two gauges becomes large

<sup>5</sup>It is written as  $(1/2)x^2$  in Ref. [14], but an overall factor was missing.

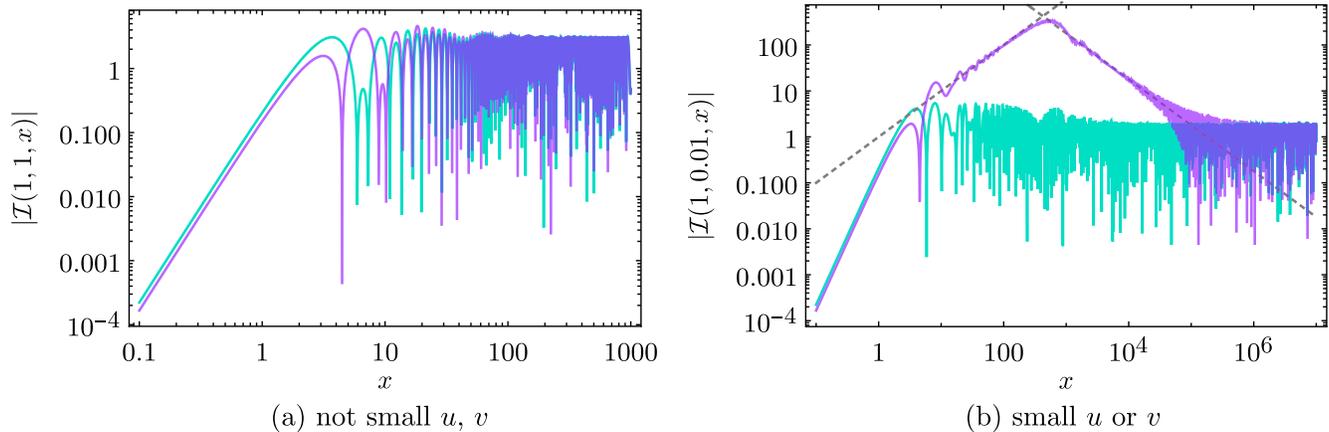


FIG. 1. Comparison of time evolution of the function  $|\mathcal{I}(u, v, x)|$  in the Newtonian gauge (cyan) and in the TT gauge (purple). The values of  $u$  and  $v$  are set as  $(u = 1, v = 1)$  in the left panel, and  $(u = 1, v = 0.01)$  in the right panel. The gray dotted lines are  $x$  and  $18/(u^2 v^2 x)$ .

for small  $u$  or  $v$  (right panel of Fig. 1). This should come from the gauge dependence of the part of the tensor perturbations coupling with scalar perturbations. After a while, the difference of  $\mathcal{I}$  decreases slowly as  $\propto x^{-1}$ . Finally, the difference becomes smaller than the asymptotic value at a late time  $x \gg 1$ , which means the tensor perturbations in the TT gauge are dominated by the freely propagating tensor perturbations (GWs). Let us see these behaviors quantitatively below.

Typically, the Newtonian gauge results (cyan line) reach the asymptotic value as early as its horizon entry (both in the left and right panel). TT gauge results (purple line) are similar if  $u$  and  $v$  are not smaller than 1 (left panel), but this changes if  $u$  or  $v$  are small, i.e.,  $uv \lesssim 1$  (right panel).<sup>6</sup> In this case,  $\mathcal{I}$  in the TT gauge once increases as  $x$  for  $3\sqrt{2} \lesssim x \lesssim 3\sqrt{2}/(uv)$  and then decreases as  $18/(u^2 v^2 x)$  for  $x \gtrsim 3\sqrt{2}/(uv)$  until it reaches the asymptotic value, whose  $(u, v)$  dependence is not simple [see Eq. (8)]. The peak of  $\mathcal{I}$ , given as  $x \simeq 3\sqrt{2}/(uv)$ , roughly corresponds to the horizon entry of the scalar perturbation with the wave number  $\min[uk, vk]$ . This means that once both the scalar perturbations with their wave numbers being  $uk$  and  $vk$  enter the horizon, the gauge dependence of the tensor perturbations starts to decrease and finally becomes negligible. This behavior is similar to the decay of the gauge dependence for the energy density perturbation on subhorizon scales [44]. In summary, the gauge dependence vanishes well after the tensor mode and the scalar modes enter the horizon. Note again the difference between the tensor perturbations in the two gauges at a finite time is due to the part of those coupling with the scalar perturbations.

<sup>6</sup>Note that  $u$  and  $v$  cannot be small simultaneously because of momentum conservation: see the integration region in Eq. (3). Also,  $uv \ll 1$  is possible only at the corners of the integration region, so the net effect of gauge dependence after integration over  $u$  and  $v$  is suppressed.

### III. DISCUSSION

In this paper, we reconsider the results in Ref. [41] and show that the late-time (e.g., observable) GWs calculated in the TT gauge and in the Newtonian gauge coincide with each other, in contrast to Ref. [41]. Here, the “late time” limit means the subhorizon limit: time much after the horizon entry of the relevant scalar source modes as well as the tensor mode. In addition, the calculation in the uniform curvature gauge (flat gauge) coincides with the Newtonian-gauge result too [40]. The main reason for the agreement in these gauges is because the tensor perturbations are finally dominated by the freely propagating tensor perturbations (GWs) in all of the three gauges.

On the other hand, the tensor perturbations could be different in the comoving gauge even in the subhorizon limit during a RD era, according to Ref. [40]. We can find that the difference between the Newtonian gauge and the comoving gauge comes from the terms whose oscillations are of the form  $\cos((u \pm v)k\eta/\sqrt{3})$  or  $\sin((u \pm v)k\eta/\sqrt{3})$ , which are not the oscillations of *gravitational waves*,  $\cos(k\eta)$  or  $\sin(k\eta)$ ; see Eq. (65) in Ref. [40]. This indicates that the gauge dependence appears only in the tensor perturbations coupling with scalar perturbation, not in the gravitational wave.<sup>7</sup>

We also comment on the gauge dependence from the viewpoint of the diffusion damping [45], which occurs on small scales, since many authors focus on the GWs induced by the small-scale scalar perturbations (often in the context of PBHs). After inducing GWs, the

<sup>7</sup>A corollary of this discussion is that the large enhancement of the induced GWs found in Ref. [19], which one may suspect a gauge artifact, is not a gauge artifact. They are produced in a RD era, just after a sudden reheating transition, and the late-time time dependence of  $I(u, v, x)$  is that of genuine gravitational waves, i.e.,  $\cos(k\eta)/a$  or  $\sin(k\eta)/a$ .

small-scale perturbations finally decay exponentially,  $\propto \exp(-k^2/k_D^2(t))$ , within the diffusion scale  $k_D^{-1}(t)$  because the free-streaming length of some species, such as neutrinos or photons, increases as the temperature decreases (see, e.g., Ref. [46] for the time dependence of the damping scale). This means that the scalar perturbations inducing the GWs during the RD era get exponentially suppressed in most cases, and therefore the gauge dependence of the tensor perturbations would disappear. Although we leave the analysis of the diffusion damping effect on the evolutions of the tensor perturbations for future work, we expect that the gauge dependence that appears in the comoving gauge [40] would disappear after the exponential damping of the scalar perturbations. This is because if the scalar perturbations are exponentially suppressed and can be negligible, the difference between the tensor perturbations in two different gauges, which should be written in terms of the scalar perturbations, can also be negligible. In summary, gauge dependence disappears due to the diffusion damping, so the late-time GWs should be unique (the same as the one calculated in the Newtonian gauge).

Note that the situation is different for the tensor perturbation induced by the scalar perturbations during the late MD era. In this case, the scalar perturbations continue to couple with the tensor perturbations on subhorizon scales even in the Newtonian gauge because of the growing matter perturbations. Therefore, the tensor perturbations are dominated by the ones coupling with the scalar perturbations. That is, this kind of tensor perturbations can be larger than the GWs induced during the RD era on large scales [11]. This is why the tensor perturbation easily depends on the gauge during a MD era. For example, in Appendix B 2, we show that the function  $I$  in the TT gauge is different from that in the Newtonian gauge even at late time. However, we should notice that this kind of tensor perturbation is not a *gravitational wave*. The tensor perturbation is time-independent in the Newtonian gauge and does not behave as radiation [11,12,14]. Since most observations assume that the tensor perturbations are *gravitational waves*, the observational sensitivities for this kind of tensor perturbation will be different from those for conventional gravitational waves. Also, the observation of this tensor perturbation might require a discussion about the suitable gauge for the observation because of its gauge dependence.

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*Note added.*—The results of the latest version of Refs. [41,48] agree with ours.

### APPENDIX A: CONVENTION

We use the reduced Planck unit,  $c = \hbar = 8\pi G = 1$ , and basically follow the notations and conventions in Refs. [14,42].

The line element is parametrized as

$$ds^2 = -a^2(1 + 2\phi)d\eta^2 + 2a^2B_{,i}d\eta dx^i + a^2 \left( (1 - 2\psi)\delta_{ij} + 2E_{,ij} + \frac{1}{2}h_{ij} \right) dx^i dx^j, \quad (\text{A1})$$

where we only write the relevant perturbation variables. The quantities  $\phi$ ,  $\psi$ ,  $B$ , and  $E$  are first-order scalar perturbations, and  $h_{ij}$  is the second-order transverse traceless tensor perturbation. The comma denotes differentiation: for example,  $E_{,ij} = \partial_i \partial_j E$ . The TT gauge is defined as the gauge where  $\phi = B = 0$ . On the other hand, the Newtonian gauge is defined as  $E = B = 0$ .

The normalizations and signs of  $\psi$ ,  $E$ , and the shear potential  $\sigma = \partial E / \partial \eta$  in the TT gauge are the same as those in Ref. [41]. If we pick up the second-order component of  $h_{ij}$  [see their Eq. (A.12g)], it is also the same in the coordinate space. The convention of the Fourier transform is different, but it does not affect most parts of the calculation. When there is no anisotropic stress, the following equations are satisfied in each gauge:

$$\phi = \psi \quad (\text{Newtonian gauge}), \quad (\text{A2})$$

$$\frac{\partial \sigma}{\partial \eta} = -(\psi + 2\mathcal{H}\sigma) \quad (\text{TT gauge}). \quad (\text{A3})$$

The relation of perturbation variables in the TT gauge and in the Newtonian gauge is given by Eqs. (5.2) and (5.3) in Ref. [41]:

$$\psi^{(\text{TT gauge})}(\eta, \mathbf{k}) = \Psi + \mathcal{H}(\eta) \left( \int_0^\eta \frac{a(\bar{\eta})}{a(\eta)} \Psi(\bar{\eta}, \mathbf{k}) d\bar{\eta} - \frac{\mathcal{D}_1(\mathbf{k})}{a(\eta)} \right), \quad (\text{A4})$$

$$\sigma^{(\text{TT gauge})}(\eta, \mathbf{k}) = - \int_0^\eta \frac{a(\bar{\eta})}{a(\eta)} \Psi(\bar{\eta}, \mathbf{k}) d\bar{\eta} + \frac{\mathcal{D}_1(\mathbf{k})}{a(\eta)}, \quad (\text{A5})$$

where  $\Psi$  is a gauge invariant Bardeen potential which reduces to  $\psi$  in the Newtonian gauge, and  $\mathcal{D}_1(\mathbf{k})$  is an integration constant which does not depend on  $\eta$ . We take  $\mathcal{D}_1(\mathbf{k}) = 0$ , which corresponds to  $\mathcal{C}(\mathbf{k}) = -2\zeta(\mathbf{k})/k^2$  in Ref. [41]: for comparison, see also their footnote 6. The transfer functions of  $\psi = \psi^{(\text{TT gauge})}$  and  $\sigma = \sigma^{(\text{TT gauge})}$  during a RD era in the TT gauge are defined as [41]<sup>8</sup>

<sup>8</sup>We take the sign notation for the curvature perturbation which is opposite to that taken in Ref. [41].

$$\psi(\eta, \mathbf{k}) = -\frac{2}{3}\zeta(\mathbf{k})T_\psi(\eta, k), \quad (\text{A6})$$

$$\sigma(\eta, \mathbf{k}) = -\frac{2}{3}\zeta(\mathbf{k})\frac{\sqrt{3}}{k}T_\sigma(\eta, k), \quad (\text{A7})$$

where  $\zeta(\mathbf{k})$  is the primordial curvature perturbations.

Table I summarizes the correspondence of the notations in this paper (as well as Refs. [14,18,19,42]) and in Ref. [41]. See the last part of Appendix D of Ref. [13] for complementary information. (The notation of Ref. [41] is largely the same as that of Ref. [13].)

TABLE I. Comparison of notations. In this table, the wave-number of the tensor and scalar perturbations are denoted by  $\mathbf{k}$  and  $\mathbf{p}$ , respectively, with  $k = |\mathbf{k}|$  and  $p = |\mathbf{p}|$ .  $\bar{\eta}$  is the conformal time as an integration variable.

This paper	Ref. [41]	Correspondence
$v = p/k$	$x = p/k$	$v \leftrightarrow x$
$u =  \mathbf{k} - \mathbf{p} /k$	$y =  \mathbf{k} - \mathbf{p} /k$	$u \leftrightarrow y$
$x = k\eta$	$z = k\eta/\sqrt{3}$	$x \leftrightarrow \sqrt{3}z$
$\bar{x} = k\bar{\eta}$	$u = k\bar{\eta}$	$\bar{x} \leftrightarrow u$
$\mathcal{I} = xI$	$\mathcal{I}_c \cos x + \mathcal{I}_s \sin x$	$9\mathcal{I}_c \leftrightarrow \mathcal{I}_c,$ $9\mathcal{I}_s \leftrightarrow \mathcal{I}_s$
$= \mathcal{I}_c \cos x + \mathcal{I}_s \sin x$		

## APPENDIX B: ANALYTICAL EXPRESSION OF FUNCTION $I(u, v, x)$ IN TT GAUGE

### 1. Radiation-dominated era

In a RD era, the Eq. (4) is integrated as follows:

$$\begin{aligned}
I(u, v, x) = & \frac{3}{u^3 v^3 x^3} \left( -6uvx - uv(u^2 + v^2 - 3)x^2 \sin x + 6\sqrt{3}v \sin \frac{ux}{\sqrt{3}} + 6\sqrt{3}u \sin \frac{vx}{\sqrt{3}} - 6\sqrt{3}u \cos \frac{ux}{\sqrt{3}} \sin \frac{vx}{\sqrt{3}} \right. \\
& - 6\sqrt{3}v \cos \frac{vx}{\sqrt{3}} \sin \frac{ux}{\sqrt{3}} + 3(u^2 + v^2 - 3)x \sin \frac{ux}{\sqrt{3}} \sin \frac{vx}{\sqrt{3}} + 6uvx \cos \frac{ux}{\sqrt{3}} \cos \frac{vx}{\sqrt{3}} \left. \right) \\
& + \frac{3(u^2 + v^2 - 3)^2}{4u^3 v^3 x} \left( \sin x \left( -\text{Ci} \left( x + \frac{(u+v)x}{\sqrt{3}} \right) - \text{Ci} \left( \left| x - \frac{(u+v)x}{\sqrt{3}} \right| \right) + \text{Ci} \left( x + \frac{(u-v)x}{\sqrt{3}} \right) \right. \right. \\
& + \text{Ci} \left( x + \frac{(-u+v)x}{\sqrt{3}} \right) + \log \left| \frac{3 - (u+v)^2}{3 - (u-v)^2} \right| \left. \right) + \cos x \left( \text{Si} \left( x + \frac{(u+v)x}{\sqrt{3}} \right) + \text{Si} \left( x - \frac{(u+v)x}{\sqrt{3}} \right) \right. \\
& \left. \left. - \text{Si} \left( x + \frac{(u-v)x}{\sqrt{3}} \right) - \text{Si} \left( x + \frac{(-u+v)x}{\sqrt{3}} \right) \right) \right), \quad (\text{B1})
\end{aligned}$$

where  $\text{Si}(x) \equiv \int_0^x d\bar{x} \frac{\sin \bar{x}}{\bar{x}}$  and  $\text{Ci}(x) \equiv -\int_x^\infty d\bar{x} \frac{\cos \bar{x}}{\bar{x}}$ . The counterpart in the Newtonian gauge is Eq. (22) of Ref. [14] [see also Ref. [13]].

This formula can also be expressed in the form

$$xI(u, v, x) = \mathcal{I}_c(u, v, x) \cos x + \mathcal{I}_s(u, v, x) \sin x, \quad (\text{B2})$$

where

$$\begin{aligned}
\mathcal{I}_c(u, v, x) = & \frac{1}{u^3 v^3 x^3} \left( 9x \cos x \left( -2uvx + 2\sqrt{3}u \sin \frac{vx}{\sqrt{3}} + 2\sqrt{3}v \sin \frac{ux}{\sqrt{3}} + 2uvx \cos \frac{ux}{\sqrt{3}} \cos \frac{vx}{\sqrt{3}} \right. \right. \\
& - 2\sqrt{3}u \cos \frac{ux}{\sqrt{3}} \sin \frac{vx}{\sqrt{3}} - 2\sqrt{3}v \cos \frac{vx}{\sqrt{3}} \sin \frac{ux}{\sqrt{3}} + (u^2 + v^2 - 3)x \sin \frac{ux}{\sqrt{3}} \sin \frac{vx}{\sqrt{3}} \left. \right) \\
& - 3 \sin x \left( 6uvx \left( 1 + \cos \frac{ux}{\sqrt{3}} + \cos \frac{vx}{\sqrt{3}} \right) - 12\sqrt{3}u \sin \frac{vx}{\sqrt{3}} - 12\sqrt{3}v \sin \frac{ux}{\sqrt{3}} \right. \\
& - 18uvx \cos \frac{ux}{\sqrt{3}} \cos \frac{vx}{\sqrt{3}} + 3(u^2 + v^2 + 3)x \sin \frac{ux}{\sqrt{3}} \sin \frac{vx}{\sqrt{3}} \\
& \left. \left. + \sqrt{3}u(12 + (u^2 - v^2 - 3)x^2) \cos \frac{ux}{\sqrt{3}} \sin \frac{vx}{\sqrt{3}} + (u \leftrightarrow v) \right) \right) \\
& + \frac{3(u^2 + v^2 - 3)^2}{4u^3 v^3} \left( \text{Si} \left( x + \frac{(u+v)x}{\sqrt{3}} \right) + \text{Si} \left( x - \frac{(u+v)x}{\sqrt{3}} \right) \right. \\
& \left. - \text{Si} \left( x + \frac{(u-v)x}{\sqrt{3}} \right) - \text{Si} \left( x + \frac{(-u+v)x}{\sqrt{3}} \right) \right), \quad (\text{B3})
\end{aligned}$$

$$\begin{aligned}
\mathcal{I}_s(u, v, x) = & \frac{1}{u^3 v^3 x^3} \left( -3uv(u^2 + v^2 - 3)x^3 + 9x \sin x \left( -2uvx + 2\sqrt{3}v \sin \frac{ux}{\sqrt{3}} + 2\sqrt{3}u \sin \frac{vx}{\sqrt{3}} \right. \right. \\
& - 2\sqrt{3}u \cos \frac{ux}{\sqrt{3}} \sin \frac{vx}{\sqrt{3}} + (u \leftrightarrow v) + 2uvx \cos \frac{ux}{\sqrt{3}} \cos \frac{vx}{\sqrt{3}} + (u^2 + v^2 - 3)x \sin \frac{ux}{\sqrt{3}} \sin \frac{vx}{\sqrt{3}} \Big) \\
& + 3 \cos x \left( 6uvx \left( 1 + \cos \frac{ux}{\sqrt{3}} + \cos \frac{vx}{\sqrt{3}} \right) - 12\sqrt{3}v \sin \frac{ux}{\sqrt{3}} - 12\sqrt{3}u \sin \frac{vx}{\sqrt{3}} \right. \\
& - 18uvx \cos \frac{ux}{\sqrt{3}} \cos \frac{vx}{\sqrt{3}} + 3(u^2 + v^2 + 3)x \sin \frac{ux}{\sqrt{3}} \sin \frac{vx}{\sqrt{3}} \\
& \left. \left. + \sqrt{3}(12 + (u^2 - v^2 - 3)x^2) \cos \frac{ux}{\sqrt{3}} \sin \frac{vx}{\sqrt{3}} + (u \leftrightarrow v) \right) \right) \\
& - \frac{3(u^2 + v^2 - 3)^2}{4u^3 v^3} \left( \text{Ci} \left( x + \frac{(u+v)x}{\sqrt{3}} \right) + \text{Ci} \left( \left| x - \frac{(u+v)x}{\sqrt{3}} \right| \right) \right. \\
& \left. - \text{Ci} \left( x + \frac{(u-v)x}{\sqrt{3}} \right) - \text{Ci} \left( x + \frac{(-u+v)x}{\sqrt{3}} \right) - \log \left| \frac{3 - (u+v)^2}{3 - (u-v)^2} \right| \right), \tag{B4}
\end{aligned}$$

where  $(u \leftrightarrow v)$  denotes the term obtained by exchanging  $u$  and  $v$  in the previous term.

## 2. Matter-dominated era

In the case of tensor perturbations induced in a MD era, we take the following normalizations for the transfer functions of  $\psi$  and  $\sigma$ :

$$\psi(\eta, \mathbf{k}) = -\frac{3}{5} \zeta(\mathbf{k}) T_\psi(\eta, k), \tag{B5}$$

$$\sigma(\eta, \mathbf{k}) = -\frac{3}{5} \zeta(\mathbf{k}) \frac{\sqrt{3}}{k} T_\sigma(\eta, k). \tag{B6}$$

Then, the concrete expressions of the transfer functions are given as [41]

$$T_\psi(x) = \frac{5}{3}, \quad T_\sigma(x) = \frac{-x}{3\sqrt{3}}. \tag{B7}$$

Taking into account the relation  $\Psi = -(3 + 3w)/(5 + 3w)\zeta$  on the superhorizon scales where  $w = P/\rho$  is the equation-of-state parameter with  $P$  denoting the pressure, we define  $f(u, v, x)$  in a MD era  $(3/5)^2(2/3)^{-2} (= 81/100)$  times Eq. (5). Strictly speaking, the coefficient of the last term should also be modified. However,  $T'_\psi$  vanishes in a MD era, so it does not matter. The explicit form of  $f$  is  $f(u, v, x) = 1 - x^2/50$ , which grows in time.

The function  $I(u, v, x)$  in the MD era is

$$I(u, v, x) = \frac{-x^5 + 60x^3 + 180x \cos x - 180 \sin x}{50x^3}. \tag{B8}$$

The late-time limit is  $I(u, v, x \rightarrow \infty) = -x^2/50$ . Note that the counterpart in the Newtonian gauge is  $6/5$  [Eq. (36) of Ref. [14]; see also Ref. [12]].

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